

Moduli Decays

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At $H \simeq M_\Phi$: Φ field starts to oscillate coherently; same as collection of Φ particles at rest.

$$\rho_\Phi \propto a^{-3}, \quad \rho_{\text{rad}} \propto a^{-4} \implies$$

If lifetime τ_Φ is sufficiently large, ρ_Φ will dominate the total energy density: “period of early matter domination”!

Moduli Cosmology (cont'd)

Typically: Φ couplings suppressed by $1/M_{\text{Pl}} \implies$

$$\Gamma_{\Phi} = \kappa \frac{M_{\Phi}^3}{M_{\text{Pl}}^2} \quad \kappa = \frac{C}{8\pi}, \quad M_{\text{Pl}} = 1.22 \cdot 10^{19} \text{ GeV}$$

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Define $T_{\text{RH}} = \sqrt{\Gamma_{\Phi} M_{\text{Pl}}} \left(\frac{45}{4\pi^3 g_*(T_{\text{RH}})} \right)^{1/4}$

g_* : no. of d.o.f. in radiation

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$$\text{BBN} \implies T_{\text{RH}} > \text{few GeV} \implies \\ M_{\Phi} > 10^5 \kappa^{-1/3} \text{ GeV}$$

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- Direct $\Phi \rightarrow \chi$ decays: increase Ω_χ
- ρ_Φ increases $H \implies$ increases Ω_χ (if χ was in equilibrium)
- Φ decays increase entropy: reduces Ω_χ (after normalization to CMB)

Equations

$$H^2 = \frac{8\pi\rho_{\text{tot}}}{3M_{\text{Pl}}^2} \quad \rho_{\text{tot}} = \rho_{\text{rad}} + \rho_{\chi} + \rho_{\Phi}$$

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$$\frac{ds}{dt} + 3Hs \simeq \frac{1}{T}(1 - \bar{B})\Gamma_{\Phi}\rho_{\Phi}$$

$$\rho_{\text{rad}} = \frac{\pi^2}{30}g_*T^4, \quad s = \frac{2\pi^2}{45}h_*T^3 : \text{entropy density}$$

$$B_{\chi} \propto \text{Br for } \Phi \rightarrow \chi \text{ decays}$$

Note: This form of eqs works even if g_*, h_* are not constant!
They change rapidly around $T = 150 \text{ MeV}$.

Results

MD, F. Hajkarim, JCAP 1802 (2018), 057 & JHEP 1812 (2018) 042

Generally: final Ω_χ depends on $\langle\sigma v\rangle$, m_χ , B_χ , Γ_Φ , M_Φ , initial densities

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If $\rho_\Phi \gg \rho_{\text{rad}}, \rho_\chi$ at some “initial” time: B_χ only enters as ratio B_χ/M_Φ , e.g.

$$\Omega_\chi^{\Phi\text{-decay}} \propto \frac{T_{\text{RH}} B_\chi}{M_\Phi}$$

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One application:

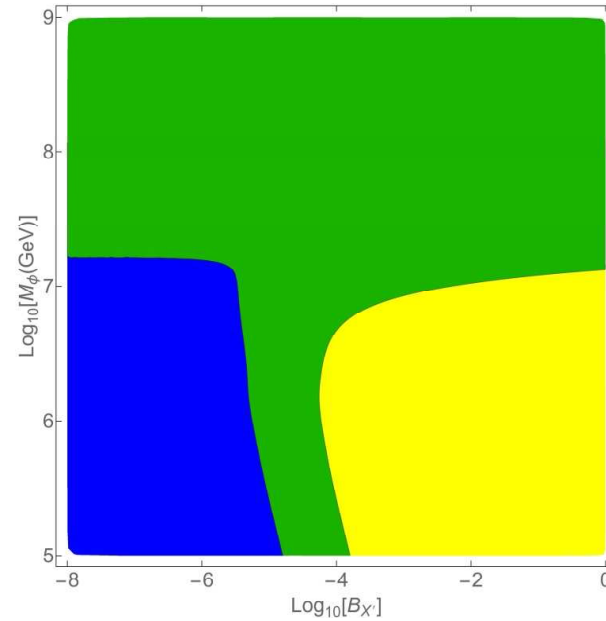
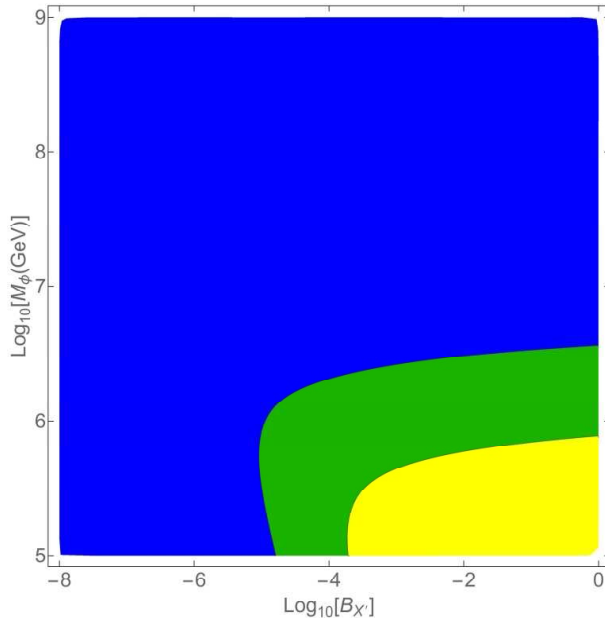
WIMP with $\langle\sigma v\rangle < \langle\sigma v\rangle_{\text{thermal}}$ WIMP: has too high relic density in standard cosmology.

Example: bino-like LSP!

Can be ok if $B_\chi < 10^{-4}$ and $M_\Phi \lesssim 10^7$ GeV! ($\kappa = 1$)

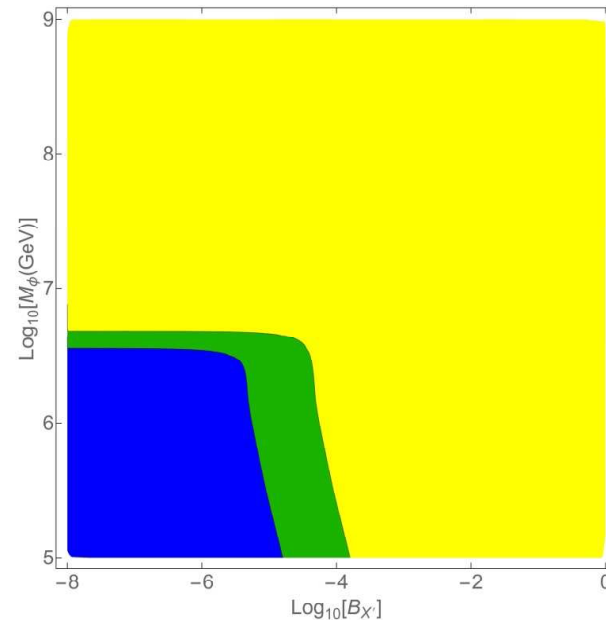
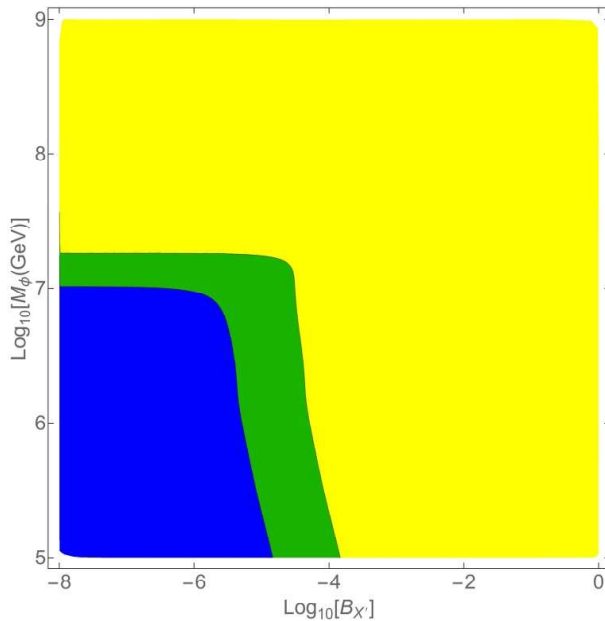
$$m_\chi = 100 \text{ GeV}$$

$\sigma_{\text{ann}} =$
 10^{-6} GeV^{-2}
 (“too large”)



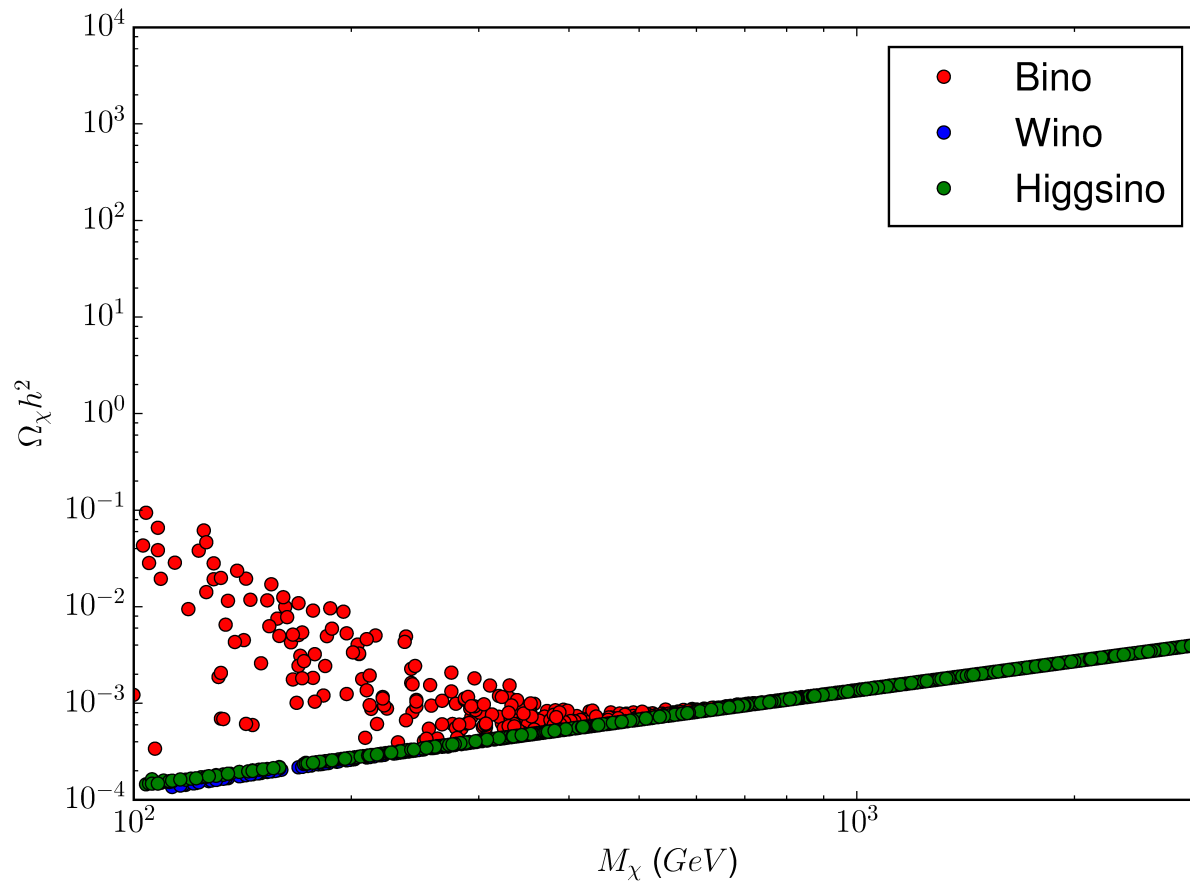
10^{-8} GeV^{-2}
 (“just right”)

10^{-9} GeV^{-2}
 (“too small”)



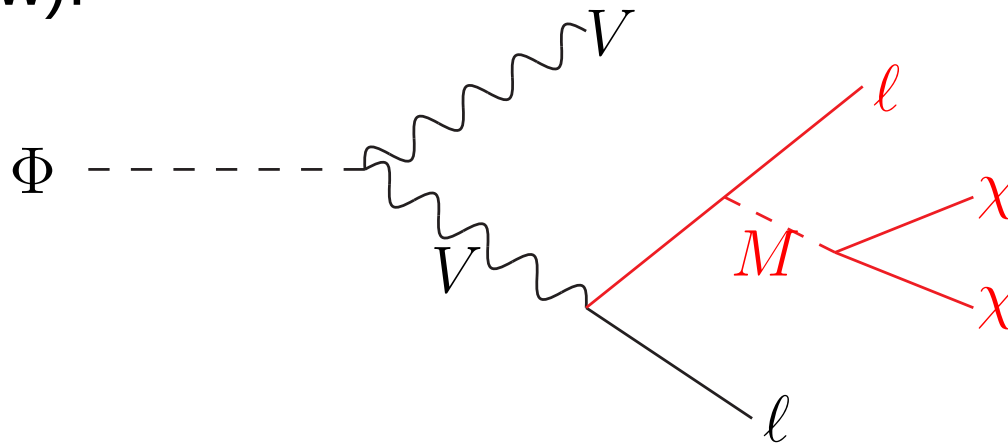
$10^{-14} \text{ GeV}^{-2}$
 (“much too small”)

In the pMSSM: $B_\chi = 10^{-5}$, $M_\Phi = 5 \cdot 10^6$ GeV



Is that realistic?

In “simplified model”: No Problem! E.g. couple WIMP χ via mediator M to leptons; Φ decays into gauge bosons (see below):



Is 5-body final state: $B_\chi \sim \alpha \alpha_\chi^2$ is automatically small if $\langle \sigma v \rangle$ is small due to small coupling α_χ .

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$$\implies B_\chi = 2 \sum_i \Gamma(\Phi \rightarrow \text{sparticle}_i) / \Gamma_\Phi^{\text{tot}}$$

$B_\chi \ll 1$ only possible if *all* $\Phi \rightarrow \text{sparticle}$ decays are strongly suppressed!

At leading order

Assumption: Φ is SM gauge singlet! Gauge invariant couplings:

- **dim. 3:** $\Phi H_u \cdot H_d$: “Higgs” mode
- **dim. 5:** $\Phi G_{\mu\nu} G^{\mu\nu}$, $\Phi W_{\mu\nu} W^{\mu\nu}$, $\Phi B_{\mu\nu} B^{\mu\nu}$: “gauge”
 $\Phi \overline{f_L} f_R H$: “fermion” mode

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Can these vertices be derived from supersymmetric Lagrangians?

Let $\hat{\Phi} = (\Phi, \tilde{\Phi})$: chiral modulus superfield.

Higgs mode

Superpot. $W \supset \hat{\Phi} \hat{H}_u \cdot \hat{H}_d$: SUSY invariant, but would give
 $\Gamma(\Phi \rightarrow H_u H_d) \ll \Gamma(\Phi \rightarrow \tilde{h}_u \tilde{h}_d)$
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Instead: introduce SUSY breaking
spurion $\hat{F} = \theta\theta F$

$$\Rightarrow \mathcal{L} \supset \frac{c_H}{M_{\text{Pl}}} \int d^2\theta \hat{F} \hat{\Phi} \hat{H}_u \cdot \hat{H}_d$$

$\Rightarrow \Phi H_u H_d$ **vertex factor** $c_H F/M_{\text{Pl}}$: is soft breaking term!

Fermion mode

$\overline{f_L} f_R H$ comes from $W \supset \hat{f}_L \hat{f}_R^c \hat{H}$:

Also generates $\overline{f_L} \tilde{f}_R \tilde{h}$, $\tilde{f}_L \overline{f_R} \tilde{h}$ terms

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$\Rightarrow$ supersymmetric $\Phi f \bar{f} H$ vertex will *always* lead to $B_\chi \sim \mathcal{O}(1)$, if $M_\Phi \gg m_{\text{particle}}$!

Gauge mode

$$\mathcal{L} \supset \frac{c_g}{M_{\text{Pl}}} \int d^2\theta \hat{\Phi} \hat{W}_A \hat{W}^A$$

$\hat{W}_A$: Gauge field strength superfield

$\Rightarrow \Phi$ couples to $\int d^2\theta \hat{W}_A \hat{W}^A = -\frac{1}{4} G_{\mu\nu}^a G^{\mu\nu a} + i\lambda^a \sigma^\mu D_\mu \lambda^a$

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Also gives rise to $\Phi \tilde{g} \tilde{g}$ vertex, but

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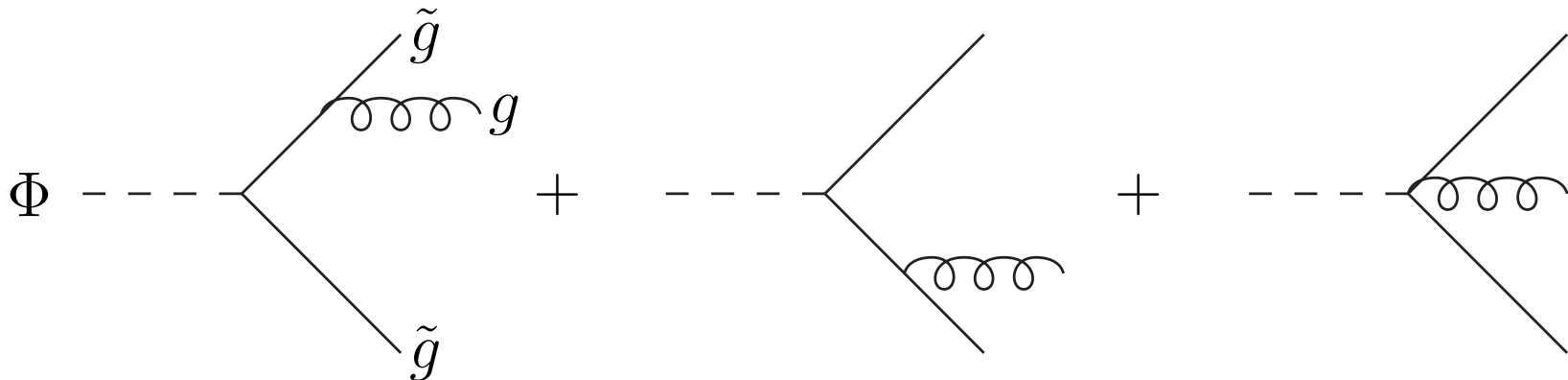
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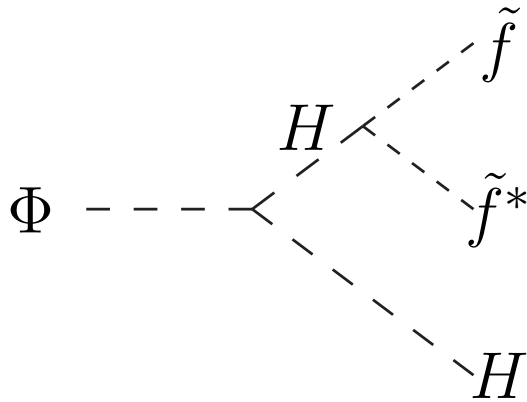
$$\mathcal{M}(\Phi \rightarrow \tilde{g} \tilde{g}) \propto m_{\tilde{g}} \text{ after using e.o.m.! Need } m_{\tilde{g}} \lesssim 10^{-2} M_\Phi.$$

Remains true for $\mathcal{M}(\Phi \rightarrow \tilde{g} \tilde{g} g)$ from:



Higher Orders: Higgs

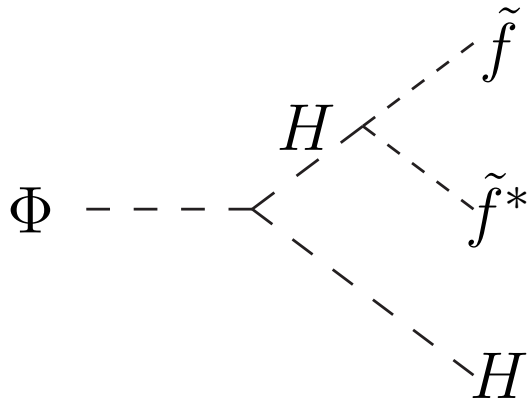
A. Banik, MD, to appear



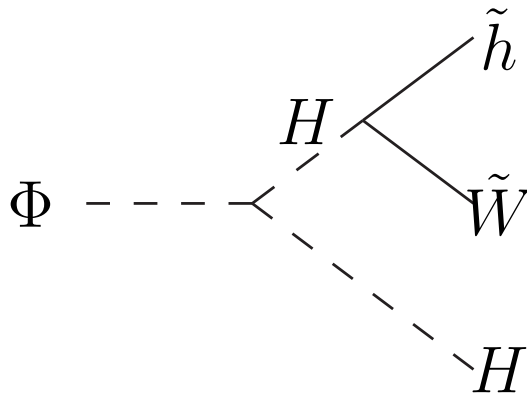
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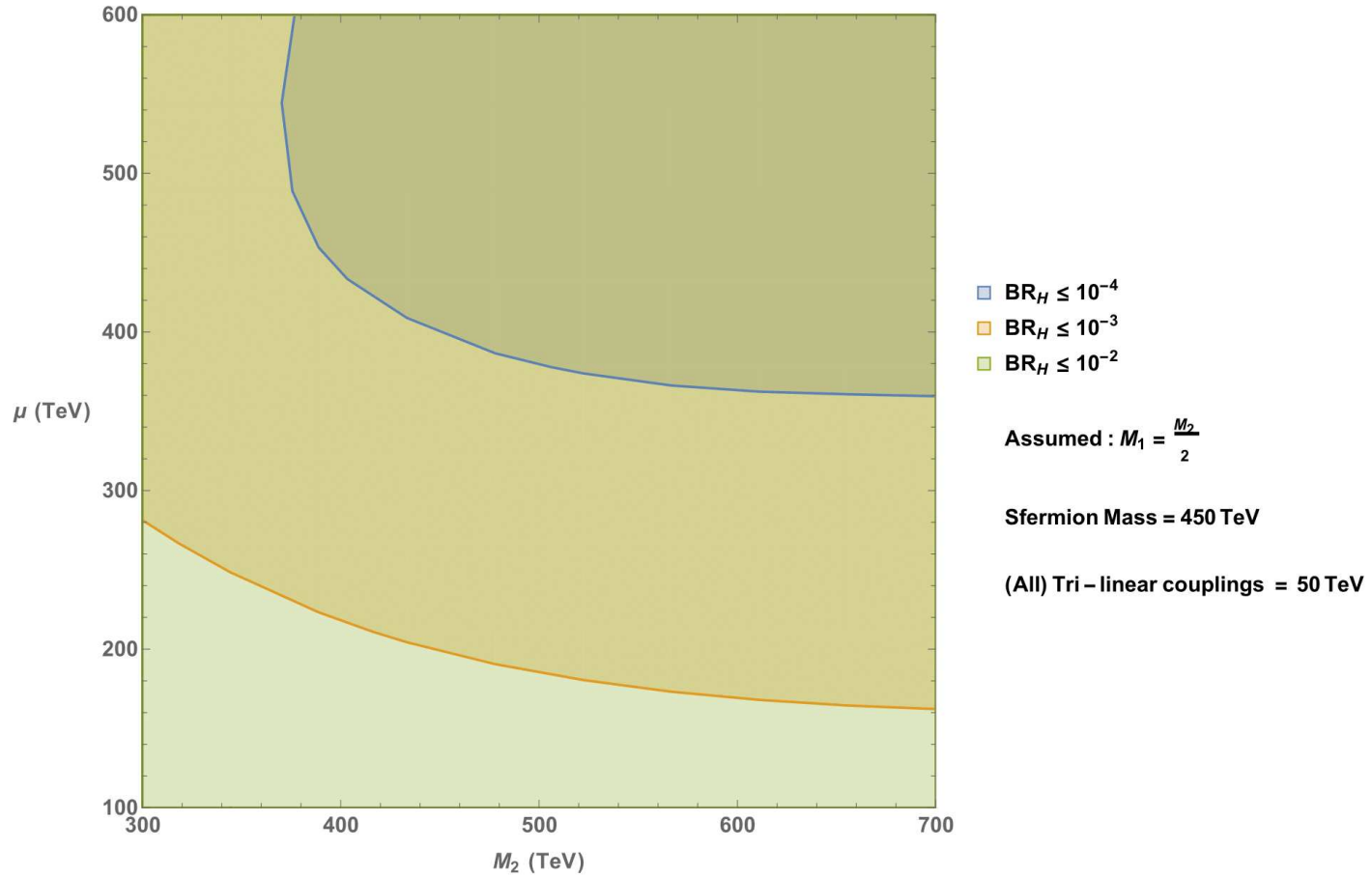


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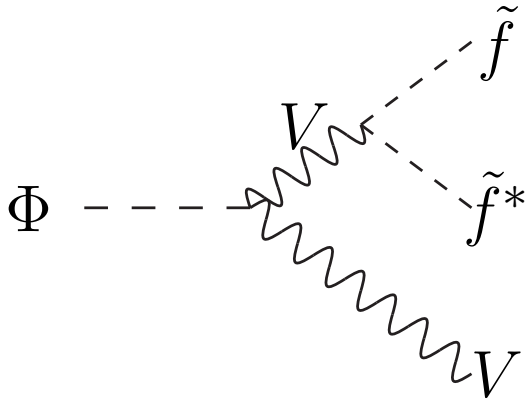


No extra suppression,
except elw cplg, phase space!

Results: $M_\Phi = 10^6 \text{ GeV}$, $\Phi \rightarrow H_u H_d$

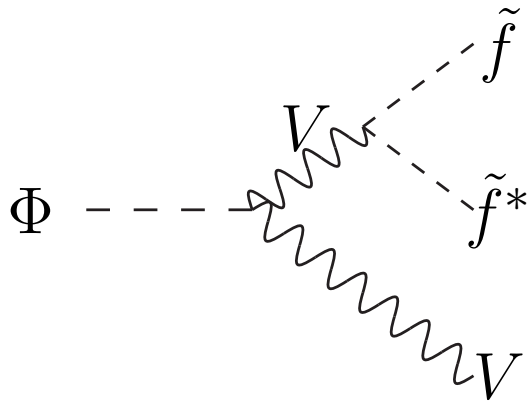


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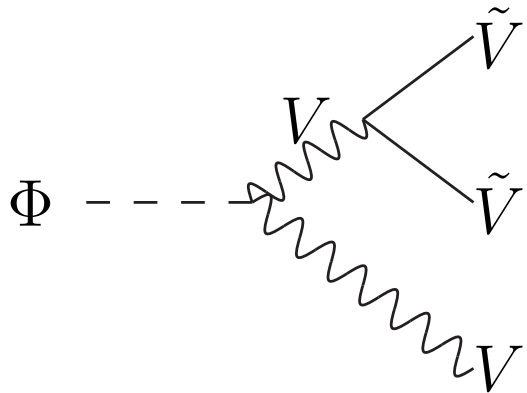


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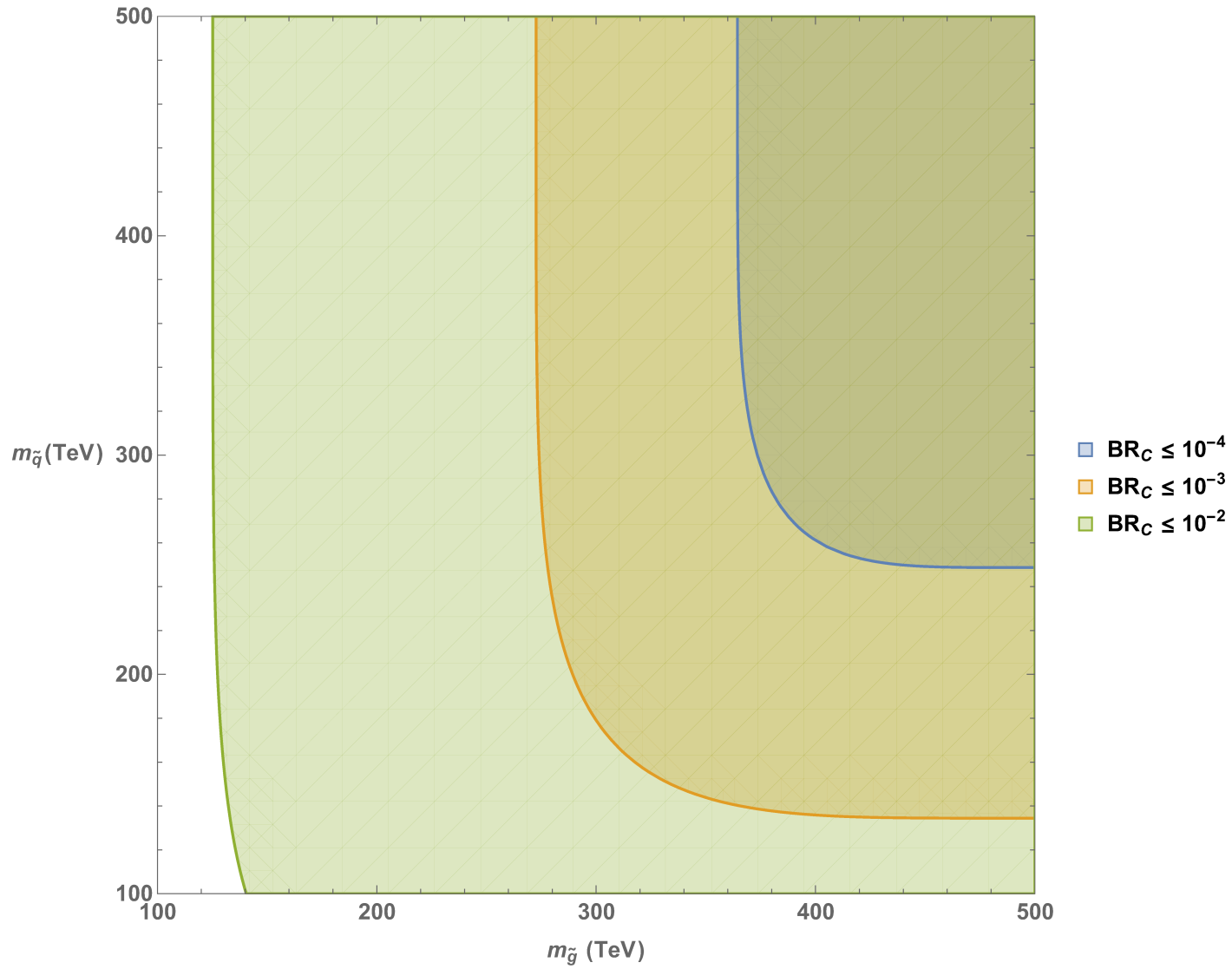


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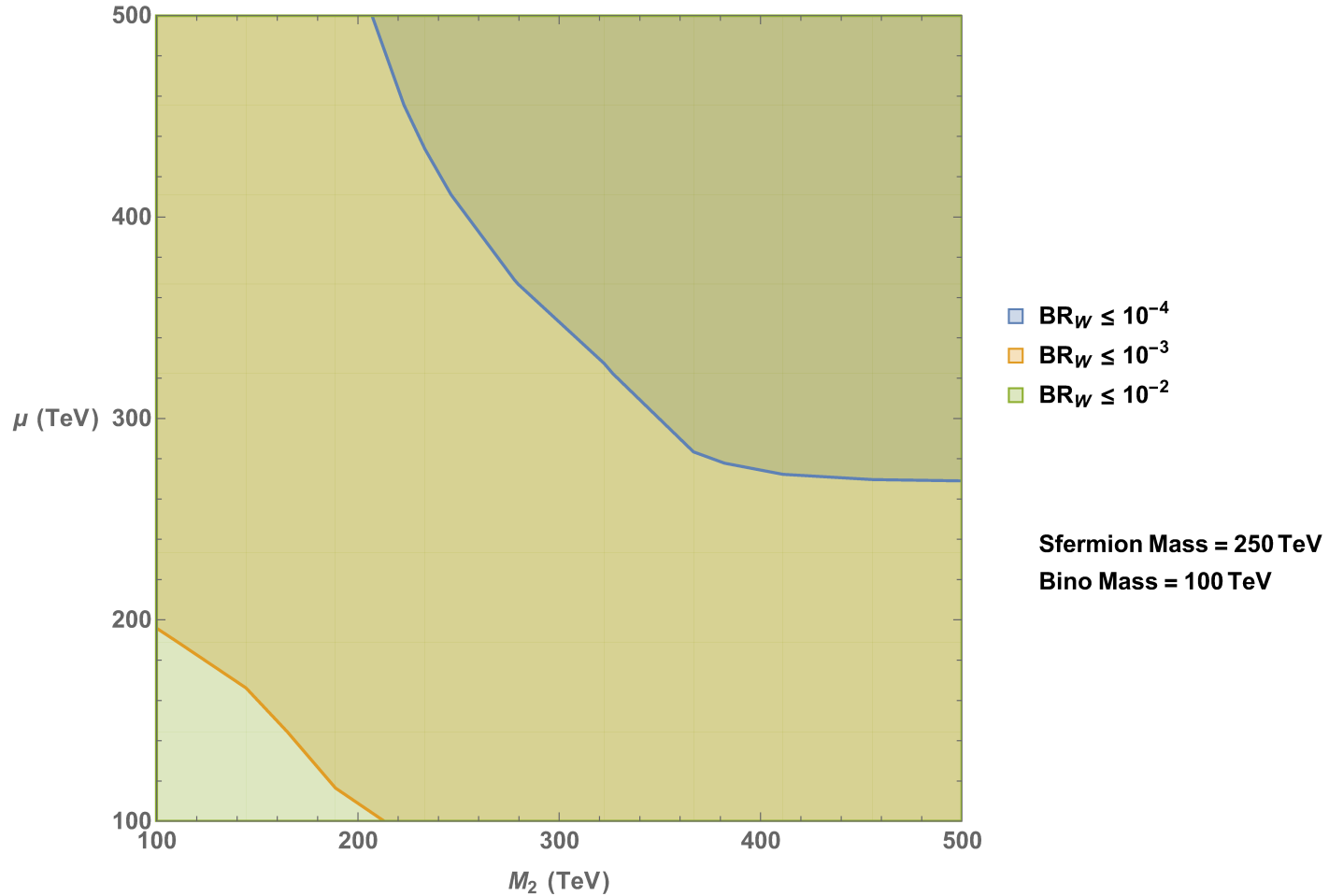


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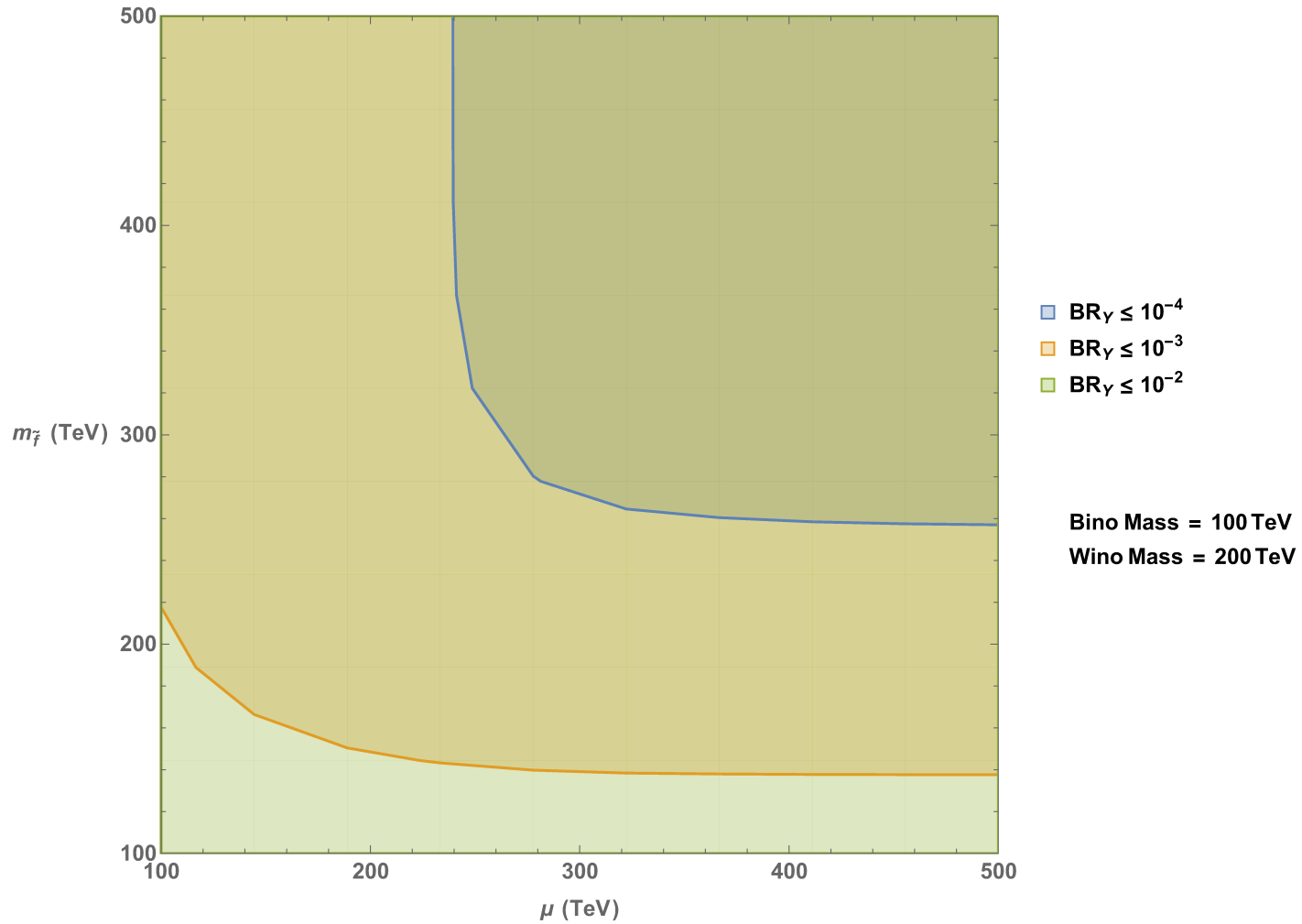
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Results: $M_\Phi = 10^6$ GeV, $\Phi \rightarrow BB$



Higher order decays: Summary

$B_\chi \gtrsim 10^{-2}$ (10^{-3}) for LO Φ decays into strongly (weakly) interacting particles, **unless** $m_{\text{particle}} \gtrsim M_\Phi/3$: **Not natural!**

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- Higgs mode: $\Phi H_u H_d$ cplg should be much smaller than usual soft breaking masses.

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- At leading order: can write supersymmetric $\Phi gg/\Phi VV$ vertex, soft breaking $\Phi H_u H_d$ vertex with desired properties
- Three-body final states will give $B(\Phi \rightarrow \tilde{\chi}_1^0) > 10^{-3}$ if $m_{\text{particle}} \ll M_\Phi$; can be corrected by increasing M_Φ , if Φ couplings are $\lesssim 0.03/M_{\text{Pl}}$.